

COMPARING THE DARK MATTER HALOS OF SPIRAL, LOW SURFACE BRIGHTNESS AND DWARF SPHEROIDAL GALAXIES

MATTHEW G. WALKER¹, STACY S. MCGAUGH², MARIO MATEO³, EDWARD W. OLSZEWSKI⁴ AND RACHEL KUZIO DE NARAY^{5,6}

Draft version May 17, 2010

ABSTRACT

We consider dark masses measured from kinematic tracers at discrete radii in galaxies for which baryonic contributions to overall potentials are either subtracted or negligible. Recent work indicates that rotation curves due to dark matter (DM) halos at intermediate radii in spiral galaxies are remarkably similar, with a mean rotation curve given by $\log_{10}[V_{c,DM}/(\text{kms}^{-1})] = 1.47^{+0.15}_{-0.19} + 0.5 \log_{10}[r/\text{kpc}]$. Independent studies show that while estimates of the dark mass of a given dwarf spheroidal (dSph) galaxy are robust only near the half-light radius, data from the Milky Way's (MW's) dSph satellites are consistent with a narrow range of mass profiles. Here we combine published constraints on the dark halo masses of spirals and dSphs and include available measurements of low surface brightness galaxies for additional comparison. We find that most measured MW dSphs lie on the extrapolation of the mean rotation curve due to DM in spirals. The union of MW-dSph and spiral data appears to follow a mass-radius relation of the form $M_{DM}(r)/M_{\odot} = 200^{+200}_{-120}(r/\text{pc})^2$, or equivalently a constant acceleration $g_{DM} = 3^{+3}_{-2} \times 10^{-9} \text{cms}^{-2}$, spanning $0.02 \lesssim r \lesssim 75 \text{ kpc}$. Evaluation at specific radii immediately generates two results from the recent literature: a common mass for MW dSphs at fixed radius and a constant DM central surface density for galaxies ranging from MW dSphs to spirals. However, recent kinematic measurements indicate that M31's dSph satellites are systematically less massive than MW dSphs of similar size. Such deviations from what is otherwise a surprisingly uniform halo relation presumably hold clues to individual formation and evolutionary histories.

Subject headings: galaxies: kinematics and dynamics — galaxies: spiral — galaxies: dwarf — galaxies: fundamental parameters — (cosmology:) dark matter

1. INTRODUCTION

Considered within the framework of Newtonian mechanics, the observed motions of stars and gas imply that the luminous components of galaxies are embedded within extended halos of unseen material. Absent the ability to measure invisible masses directly, for a given galaxy one measures dynamical and luminous masses independently and then identifies dark matter simply as the difference between the two: $M_{DM} = M_{dyn} - M_{lum}$. Here we shall investigate the properties of dark matter halos only insofar as they are accessible via the direct application of this equation, without adopting any particular halo model.

This task is made difficult by the fact that there are no galaxies for which both M_{dyn} and M_{lum} are easily measured. While the circular motions of stars and gas in spiral galaxies provide a relatively clean measure of the dynamical mass profile, subtraction of the luminous contribution suffers from systematic uncertainties, primarily those related to stellar mass-to-light ratios. One sidesteps this problem, or at least exchanges it for others, by considering dwarf spheroidal (dSph) galaxies, for which baryonic components contribute negligibly to internal gravitational potentials. However, because dSphs are supported by stellar velocity dispersions instead of rotation, estimation of their individual mass profiles requires large kinematic data sets and typically employs strong modeling assumptions (see discussion by, e.g., Pryor & Kormendy 1990;

Wilkinson et al. 2002; Gilmore et al. 2007). Larger elliptical galaxies (see Napolitano et al. 2010) combine the most challenging aspects of both regimes, as their significant stellar masses are supported in large part by random motions.

Nevertheless, recent work indicates that dark masses are constrained reasonably well at intermediate radii of both spiral and dSph galaxies. McGaugh et al. (2007, “M07”) use rotation curve data for a sample of 60 spirals and discard data points interior to $r < 1 \text{ kpc}$, where dynamical masses can be affected by noncircular motions. After subtracting baryonic masses, M07 find that the rotation curves due exclusively to dark matter halos lie nearly on top of each other, with mean rotation curve

$$\log_{10}[V_{c,DM}/(\text{kms}^{-1})] = 1.47^{+0.15}_{-0.19} + 0.5 \log_{10}[r/\text{kpc}]. \quad (1)$$

Working at smaller scales, Walker et al. (2009, “W09”) use kinematic data for eight Milky Way (MW) dSphs to show that for a wide range of plausible halo shapes and velocity anisotropies, the estimated mass within the projected half-light radius is approximately (subject to validity of the assumptions of spherical symmetry, dynamical equilibrium, flat velocity dispersion profiles $\langle v^2 \rangle^{1/2}(R) = \sigma$, and negligible contributions to measured velocity dispersions from unresolved binary motions)

$$M(r_{\text{half}}) \approx \frac{5r_{\text{half}}\sigma^2}{2G}. \quad (2)$$

Subsequently Wolf et al. (2009) have provided an analytic argument for why such an estimate is insensitive to anisotropy. Since dSph kinematics tend to be dominated by dark matter at all radii (Section 2.3; see also Mateo 1998 and references therein; Walker et al. 2007), Equation 2 provides a crude estimate of the dark mass within r_{half} for any dSph for which

walker@ast.cam.ac.uk

¹ Institute of Astronomy, University of Cambridge, UK

² Department of Astronomy, University of Maryland, College Park, MD

³ Department of Astronomy, University of Michigan, Ann Arbor

⁴ Steward Observatory, The University of Arizona, Tucson, AZ

⁵ Center for Cosmology, University of California, Irvine

⁶ NSF Astronomy & Astrophysics Postdoctoral Fellow

measurements of r_{half} and σ are available. W09 apply Equation 2 to data for 28 classical dSphs and ultrafaint satellites in the Local Group and find that, allowing for scatter by a factor of less than two in normalization, the ensemble of dSph data is consistent with a mass profile of the form $M \propto r^{1.4 \pm 0.4}$.

Thus M07 and W09 independently note the apparent self-similarity of spiral and dSph dark matter halo profiles, respectively. Here we combine data from those studies in order to examine whether the mean rotation curve attributable to dark matter in spiral galaxies extends to the small radii characteristic of dSphs. For further comparison we also include recent measurements by Kuzio de Naray et al. (2008) of rotation curves in low surface brightness (LSB) galaxies. The combined data set samples dark matter halos over the range $0.2 \lesssim r \lesssim 75$ kpc.

2. DATA

2.1. Spirals

The M07 data consist of HI rotation curves compiled by Sanders & McGaugh (2002, observational references therein), trimmed for quality to a formal accuracy of 5% or better (McGaugh 2005), and restricted to radii larger than $r \geq 1$ kpc because systematic uncertainties (e.g., noncircular motions, beam smearing) dominate the deconvolution of rotation curves at smaller radii. The final sample contains 686 independent, resolved rotation velocity measurements for 60 galaxies, spanning radii $1 \leq r \leq 75$ kpc. This sample covers virtually the entire range of spiral properties, including circular velocities $50 \leq V_f \leq 300$ km s⁻¹, scale radii $0.5 \leq r_d \leq 13$ kpc, baryonic masses $3 \times 10^8 < M_b < 4 \times 10^{11} M_\odot$, central surface brightnesses $19.6 \leq \mu_{0,b} \leq 24.2$ mag arcsec², and gas fractions $0.07 \leq f_g \leq 0.95$.

In order to isolate the contribution of dark matter to an observed rotation curve, M07 subtract contributions from stellar (bulge, disk) and HI components. For this task M07 use GIPSY software (van der Hulst et al. 1992) to convert numerically the observed luminous distributions into rotation curves. The most critical step in this procedure is the estimation of the stellar mass-to-light ratio, $\Upsilon_* \equiv M_*/L_*$. For each spiral, M07 adopt the value of Υ_* obtained by fitting the rotation curve with modified Newtonian dynamics (MoND; Milgrom 1983; Sanders & McGaugh 2002). MoND generally provides good fits to the rotation curves of spirals with Υ_* as the only free parameter, and the resulting estimates of Υ_* stand in excellent agreement with population synthesis models (Bell & de Jong 2001; McGaugh 2004). Note that despite this use of MoND as an empirically justified fitting formula with which to estimate Υ_* , the ensuing dynamical analysis of M07 is purely Newtonian. Alternative prescriptions for estimating Υ_* (e.g., maximum-disk or the population-synthesis models of Bell et al. 2003) give similar results for the dark matter rotation curves, albeit with more scatter (M07).

2.2. Low Surface Brightness Galaxies

Kuzio de Naray et al. (2006, 2008, 2009) report high resolution observations of LSB disk galaxies. These data are two-dimensional velocity fields of the inner regions of LSBs, augmented by long-slit and, at larger radii, HI data. Here we include data for the nine LSBs that have mass models presented in the second part of Table 3 from Kuzio de Naray et al. (2008). LSBs are dark-matter dominated down to small radii, but their stellar components are not completely negligible in the inner kpc. For consistency with the spiral data, we

adopt the MoND estimates of Υ_* from de Blok & McGaugh (1998) where available. These are not available in the cases of NGC 4395, UGC 4325, and DDO 64, for which we adopt population synthesis estimates given by Kuzio de Naray et al. (2008). The MoND estimates of Υ_* are consistent with those of population synthesis models (McGaugh 2004) with most of the velocity attributed to dark matter in either case, so it makes little difference which is adopted. The observed velocities have also been corrected for pressure support (typically ~ 8 km s⁻¹); this correction is barely perceptible in most cases. We relax the accuracy requirement to 30% in velocity for the LSBs in order to retain a reasonable amount of data.

2.3. Dwarf Spheroidals

We adopt the dSph data listed in Table 1 of W09 (references to original studies therein; see erratum for updated values of r_{half}). These data include projected half-light radii, global velocity dispersions and luminosities measured for 28 objects in the Local Group, including satellites of the MW and M31, and two Local Group dSphs—Cetus and Tucana—that are unassociated with any obvious parent system. To these data we add measurements for eight (six previously unmeasured) M31 satellites, adopting velocity dispersions measured by Kalirai et al. (2009) for AndI, AndII⁷, AndIII, AndVII, AndX and AndXIV, and velocity dispersions measured by Collins et al. (2009) for And IX⁸ and AndXII.

For each of the 34 objects in the final sample, we apply Equation 2 to estimate the dynamical mass, $M(r_{\text{half}})$, enclosed within the sphere having radius equal to the projected half-light radius. We adopt for all dSphs a crude stellar mass-to-light ratio of $\Upsilon_* = 1.5 \pm 1.0 M_\odot/L_{V,\odot}$ (e.g., Martin et al. 2008). The dark mass is then⁹

$$M_{\text{DM}}(r_{\text{half}}) = M(r_{\text{half}}) - \Upsilon_* \frac{L_V}{2^{3/2}}. \quad (3)$$

For comparison with the dark masses in spirals we explicitly consider for dSphs the values of $M_{\text{DM}}(r_{\text{half}})$ obtained from Equation 3, but because most dSphs are dominated by dark matter even at their centers, these values generally are indistinguishable from the dynamical masses $M(r_{\text{half}})$ estimated by W09.

We note that while the velocity dispersions of most dSphs are measured from luminous red giants, the published dispersions of five ultrafaint MW satellites—BooI, Coma, Segue-1, UMaII and Willman-1—are measured primarily from stars near the main-sequence turnoffs of these objects (Simon & Geha 2007; Martin et al. 2007; Geha et al. 2009). Since these stars have not yet expanded to become red giants, they can still exist as members of unresolved binary systems with separations and speeds that are smaller and faster, respectively, than those available to systems involving red giants. Because errorbars on the velocity dispersions of these objects do not include systematic errors due to the unknown level of contamination by binary orbital motions, masses currently estimated for the faintest satellites should be considered

⁷ For AndII, the new measurement of $\sigma = 7.3 \pm 0.8$ km s⁻¹ (Kalirai et al. 2009) replaces the value $\sigma = 9.3 \pm 2.7$ km s⁻¹ (Côté et al. 1999) listed in Table 1 of W09.

⁸ For AndIX, the new measurement of $\sigma = 3.0^{+2.8}_{-3.0}$ km s⁻¹ (Collins et al. 2009) replaces the value $\sigma = 6.8 \pm 2.5$ km s⁻¹ (Chapman et al. 2005) listed in Table 1 of W09.

⁹ For the Plummer surface brightness profiles adopted by W09, division of the total luminosity by $2^{3/2}$ gives the luminosity inside the sphere with radius equal to the projected half-light radius r_{half} .

as upper limits. See Olszewski et al. (1996); Hargreaves et al. (1996); Minor et al. (2010) for relevant discussions of the effects of binaries on estimates of dSph velocity dispersions.

3. COMPARING THE DARK MATTER HALOS OF SPIRALS, LSBs AND DSPhS

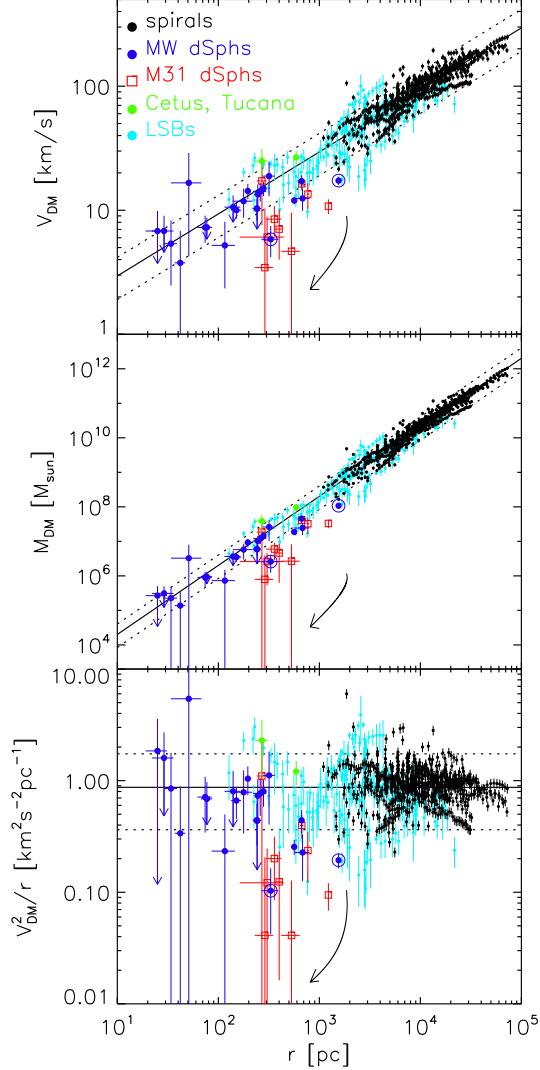


FIG. 1.— *Top*: Circular velocities due to dark matter halos at intermediate radii of spirals (data from Sanders & McGaugh 2002; McGaugh 2005; McGaugh et al. 2007), LSBs (data from Kuzio de Naray et al. 2006, 2008, 2009) and within the projected half-light radii of dwarf spheroidal galaxies (data from Walker et al. 2009 and references therein; Kalirai et al. 2009; Collins et al. 2009). Downward-pointing arrows indicate five ultrafaint dSphs (BooI, Coma, Seg1, UMaII and Will1) with velocity dispersions measured primarily from faint turnoff stars, for which unresolved binary orbital velocities may be significant. Middle and bottom panels give the same information in terms of enclosed masses and accelerations, where $g_{\text{DM}} = GM_{\text{DM}}/r^2 = V_{\text{c,DM}}^2/r$. In all panels, the solid line indicates the fit by McGaugh et al. (2007) to the spiral data alone: $\log_{10}[V_{\text{c,DM}}/(\text{kms}^{-1})] = 1.47^{+0.15}_{-0.19} + 0.5 \log_{10}[r/\text{kpc}]$, corresponding to a mass-radius relation of the form $M_{\text{DM}}(r)/M_{\odot} = 200^{+200}_{-120}(r/\text{pc})^2$, or $g_{\text{DM}} = 0.9^{+0.9}_{-0.5} \text{ km}^2 \text{ s}^{-2} \text{ pc}^{-1}$ (dotted lines indicate the formal scatter). Open circles identify the two most extreme outliers—Hercules ($r_{\text{half}} \sim 330 \text{ pc}$) and Sagittarius ($r_{\text{half}} \sim 1550 \text{ pc}$)—among MW dSph satellites. Black arrows indicate the magnitude and direction that a generic dSph would be displaced due to the tidal stripping of 99% of the original stellar mass, from N-body simulations by Peñarrubia et al. (2008).

Figure 1 displays the union of spiral, LSB and dSph data. The top panel in Figure 1 plots the rotation velocity due to the dark matter halo against the radius where this quantity is

evaluated. We re-emphasize that these estimates are derived independently of any particular (i.e., cored or cusped) halo model. Data points for spirals are the same as those in Figure 3 of M07. For the dSphs we translate estimates of $M_{\text{DM}}(r_{\text{half}})$ into circular velocities using $V_{\text{c,DM}}^2 = GM_{\text{DM}}/r$. The middle panel contains the same information cast in terms of enclosed dark mass. Here we plot dSph masses directly and convert circular velocities for spirals and LSBs into enclosed masses. In both panels, a solid line indicates the fit provided by M07 to the mean rotation curve of the spiral data alone, which is given by Equation 1 or equivalently by the mass-radius relation

$$\frac{M_{\text{DM}}(r)}{M_{\odot}} = 200^{+200}_{-120} \left(\frac{r}{\text{pc}} \right)^2. \quad (4)$$

The bottom panel of Figure 1 again conveys the same information, now cast in terms of the acceleration due to the dark matter halo, $g_{\text{DM}} = V_{\text{c,DM}}^2/r$. Since the M07 relation has the form $M \propto r^2$, the corresponding acceleration profile is constant:

$$g_{\text{DM}} = 0.9^{+0.9}_{-0.5} \text{ km}^2 \text{ s}^{-2} \text{ pc}^{-1} = 3^{+3}_{-2} \times 10^{-9} \text{ cms}^{-2}. \quad (5)$$

In order to evaluate whether the dSph and LSB data are consistent with the extrapolation of the M07 relation at small radii, we calculate

$$\chi^2 \equiv \sum_{i=1}^N \frac{(\log_{10}[V_{\text{DM},i}/(\text{kms}^{-1})] - 0.5 \log_{10}[r_i/\text{kpc}] - 1.47)^2}{\sigma_{\log_{10}[V_{\text{DM},i}/(\text{kms}^{-1})]}^2 + \sigma_{\text{M07}}^2}, \quad (6)$$

where residuals are scaled by measurement errors convolved with the intrinsic scatter in the M07 relation. The value of $\sigma_{\text{M07}} \sim 0.15$ reported by M07 reflects the non-Gaussian distribution of spiral velocities at a given radius. Here we adopt a smaller value of $\sigma_{\text{M07}} = 0.08$ so that the central 68% of spiral data points lie within $\pm \sigma_{\text{M07}}$ of the M07 relation.

For the 23 MW dSphs we obtain $\chi^2/N = 2.2$. At present the most extreme outlier is Hercules, which has a small velocity dispersion—and hence a small inferred mass—relative to other MW dSphs of similar size (Adén et al. 2009). At least some of the scatter within the dSph satellite populations likely results from tidal interactions with the MW and M31. Black arrows in Figure 1 indicate, according to N-body simulations by Peñarrubia et al. (2008), the magnitude and direction of the track followed by a generic satellite as it loses 99% of its original stellar mass to tidal disruption. Tidal evolution tends to carry individual objects toward smaller inferred masses as velocity dispersion decreases more quickly than half-light radius. Thus tides may explain the relatively small mass estimated for Hercules, whose elongated morphology (axis ratio $\sim 3 : 1$) may indicate ongoing tidal disruption (Belokurov et al. 2007; Coleman et al. 2007; Sand et al. 2009). Lending support to this explanation is the fact that the next most severe outlier among MW satellites is the main body of Sagittarius, from which tidal arms extend nearly twice around the sky (e.g., Majewski et al. 2003). If we discard Hercules and Sagittarius as systems that are strongly affected by tides, the 21 remaining MW dSphs are reasonably consistent with the M07 extrapolation, having $\chi^2/N = 1.2$. If we further exclude BooI, Com, Seg1, UMaII and Will1, we obtain $\chi^2/N = 1.5$.

The 182 points contributed by the nine LSBs have $\chi^2/N = 2.0$, indicating larger scatter than expected from observational errors and intrinsic scatter in the M07 relation. However,

since the LSBs appear to trace the M07 relation approximately, the comparison to M31 dSphs is perhaps the most compelling. Masses inferred from the new kinematic measurements of M31 dsphs (Kalirai et al. 2009; Collins et al. 2009) imply that these objects fall systematically *below* the M07 relation, with $\chi^2/N = 5.3$. Of nine measured M31 dSphs, all but AndI, AndVII and AndXV are outliers from the trend traced by spirals and MW dSphs. This behavior follows directly from the fact that at a given half-light radius the velocity dispersions measured for M31 dSphs are systematically smaller than those of MW dSphs. The M31 dSphs tend also to be larger than MW dSphs of similar luminosity (McConnachie & Irwin 2006), although this trend appears to be limited to the brightest objects (Kalirai et al. 2009). These offsets in the bulk properties of the MW and M31 satellites likely reflect differences in the environments provided by their hosts and/or the accretion histories of the satellites themselves. Whatever the reason, it is clear that in terms of the masses attributed to their dark matter halos the M31 dSphs continue to distinguish themselves from their MW counterparts.

4. DISCUSSION

M07 and W09 independently identify similarities among spiral and dSph dark mass profiles, respectively. M07 find that the rotation curves due to dark matter in spirals lie approximately on top of each other, and W09 show that the hypothesis of a “universal” dSph mass profile is as well-supported by kinematic data as independent claims of a common dSph mass scale (Mateo et al. 1993; Strigari et al. 2008, see below). Upon joining data sets, we find that spirals and MW dSphs—excluding outliers Her and Sgr and taking published velocity dispersions of the faintest satellites at face value—trace what appears to be a single relation under which dark mass scales with radius. This relation can be stated in terms of a rotation curve as in M07 (Equation 1), mass profile (Equation 4), or constant acceleration (Equation 5). Before discussing implications of the new data available for M31 dSphs, we first consider the apparent scaling of MW dSph and spiral dark matter halos in the context of other reported scaling relationships involving these objects.

If the M07 relation encodes a fundamental property of dark matter halos, then along with the appearance of a universal dSph mass profile (W09), it contains and generalizes two additional results reported in the recent literature. First, Strigari et al. (2008) model dSph dark matter halos and find that if they evaluate mass profiles at a fixed radius of 300 pc (requiring extrapolation of the mass profiles of the smallest dSphs to radii of $\sim 10r_{\text{half}}$), MW dSphs all have $M_{300} \sim 10^7 M_\odot$. If we simply evaluate the M07 mass profile at 300 pc, we obtain $M_{300} = 1.8^{+1.8}_{-1.1} \times 10^7 M_\odot$, instantly reproducing the Strigari et al. (2008) result. Here it is important to note that the M07 relation does *not* imply that all dSph dark matter halos actually extend to 300 pc, or to any other radius that may be larger than that of the observed tracer population.

Second, Kormendy & Freeman (2004), Spano et al. (2008) and Donato et al. (2009) find that if they adopt cored dark matter halos with core radius r_0 and central density ρ_0 , spiral rotation curves imply a small range in central halo surface density. Donato et al. (2009) extend this analysis to MW

dSphs and show that over 14 magnitudes in luminosity, cored dark matter halo models imply central dark matter surface densities of $r_0\rho_0 = 140^{+80}_{-30} M_\odot \text{ pc}^{-2}$. The apparent universality of $r_0\rho_0$ would then imply universality of the acceleration generated by the dark matter at r_0 , with $g_{\text{DM}}(r_0) \approx 0.5G\pi r_0\rho_0 = 3.2^{+1.8}_{-1.2} \times 10^{-9} \text{ cm s}^{-2}$ (Gentile et al. 2009). This is the same acceleration implied by the M07 relation for dark matter halos at all radii (Equation 5). Thus the M07 relation unites under a single scaling relation the appearances of a “universal mass profile” (W09) and “common mass scale” for MW dSphs (Strigari et al. 2008), and a “universal” central surface density of dark matter halos (Donato et al. 2009; Gentile et al. 2009).

The new kinematic data for M31 dSphs indicate that these objects have masses systematically smaller than their MW counterparts, placing them below the extrapolation of the M07 relation. In order to bring the M31 dSphs onto the M07 relation, either their velocity dispersions would need to be revised upward by a factor of ~ 2.5 on average or their half-light radii (perhaps involving revision of distances) would need to be revised downward to ~ 0.3 times the current estimates. Absent a compelling reason to suppose that the present M31 data are grossly affected by systematic errors, it seems that the formation and/or evolution of the MW and M31 systems have differed sufficiently to produce measurable differences in the inferred properties of the dark matter halos of their dSph satellites. For example, Peñarrubia et al. (2010) use simulations to demonstrate that tidal interactions with the baryonic disk of a host galaxy reduce the masses of dSph satellites at all radii. These simulations reproduce the observed *offset* in mass between MW and M31 dSphs by invoking a more-massive disk for M31. Note that since the degree of mass loss depends on the parent’s disk mass, this scenario requires no fine tuning of the orbital distributions of MW and M31 satellites.

Left unexplained is the puzzling circumstance that the dSph satellites of one but not both the MW and M31 appear to follow the same dark matter halo scaling relation as spiral galaxies. Here we have emphasized the consistency of MW dSph data with extrapolation of the M07 relation in order to generalize recent and independent claims of uniformity among these galaxies (Strigari et al. 2008; Donato et al. 2009; Gentile et al. 2009, W09). To the extent that the new data for M31 dSphs undermine the extrapolation of the M07 relation to dSphs, they also undermine each of these previous claims of uniformity. On the other hand, given the susceptibility of the smallest dark matter halos to evolution driven by environment, it is feasible that one of the Local Group satellite populations has evolved significantly more than the other, altering what may have been more similar conditions initially. In any case, it is likely that the emerging contrast between MW and M31 dSphs relates important details of the processes that shaped these two populations.

This work follows directly from conversations that took place at the workshop *Extreme Star Formation in Dwarf Galaxies* (Ann Arbor, MI, July 2009). We are grateful to O. Gnedin for organizing the workshop, and to J. Peñarrubia, S. Koposov and J. Wolf for helpful discussions. MGW acknowledges support from the STFC-funded Galaxy Formation and Evolution program at the Institute of Astronomy, Cambridge. SSM acknowledges support from NSF grant AST-0908370. MM and EWO acknowledge support from NSF grants AST-0808043 and AST-0807498, respectively.

REFERENCES

- Adén, D., Wilkinson, M. I., Read, J. I., Feltzing, S., Koch, A., Gilmore, G. F., Grebel, E. K., & Lundström, I. 2009, *ApJ*, 706, L150
- Bell, E. F., & de Jong, R. S. 2001, *ApJ*, 550, 212
- Bell, E. F., McIntosh, D. H., Katz, N., & Weinberg, M. D. 2003, *ApJS*, 149, 289
- Belokurov et al. 2007, *ApJ*, 654, 897
- Chapman, S. C., Ibata, R., Lewis, G. F., Ferguson, A. M. N., Irwin, M., McConnachie, A., & Tanvir, N. 2005, *ApJ*, 632, L87
- Coleman et al. 2007, *ApJ*, 668, L43
- Collins et al. 2009, *arXiv:0911.1365*
- Côté, P., Mateo, M., Olszewski, E. W., & Cook, K. H. 1999, *ApJ*, 526, 147
- de Blok, W. J. G., & McGaugh, S. S. 1998, *ApJ*, 508, 132
- Donato, F., Gentile, G., Salucci, P., Frigerio Martins, C., Wilkinson, M. I., Gilmore, G., Grebel, E. K., Koch, A., & Wyse, R. 2009, *MNRAS*, 397, 1169
- Geha, M., Willman, B., Simon, J. D., Strigari, L. E., Kirby, E. N., Law, D. R., & Strader, J. 2009, *ApJ*, 692, 1464
- Gentile, G., Famaey, B., Zhao, H., & Salucci, P. 2009, *Nature*, 461, 627
- Gilmore, G., Wilkinson, M. I., Wyse, R. F. G., Kleyna, J. T., Koch, A., Evans, N. W., & Grebel, E. K. 2007, *ApJ*, 663, 948
- Hargreaves, J. C., Gilmore, G., & Annan, J. D. 1996, *MNRAS*, 279, 108
- Kalirai et al. 2009, *arXiv:0911.1998*
- Kormendy, J., & Freeman, K. C. 2004, in *IAU Symposium*, Vol. 220, *Dark Matter in Galaxies*, ed. S. Ryder, D. Pisano, M. Walker, & K. Freeman, 377–+
- Kuzio de Naray, R., McGaugh, S. S., & de Blok, W. J. G. 2008, *ApJ*, 676, 920
- Kuzio de Naray, R., McGaugh, S. S., de Blok, W. J. G., & Bosma, A. 2006, *ApJS*, 165, 461
- Kuzio de Naray, R., McGaugh, S. S., & Mihos, J. C. 2009, *ApJ*, 692, 1321
- Majewski, S. R., Skrutskie, M. F., Weinberg, M. D., & Ostheimer, J. C. 2003, *ApJ*, 599, 1082
- Martin, N. F., de Jong, J. T. A., & Rix, H.-W. 2008, *ApJ*, 684, 1075
- Martin, N. F., Ibata, R. A., Chapman, S. C., Irwin, M., & Lewis, G. F. 2007, *MNRAS*, 380, 281
- Mateo, M., Olszewski, E. W., Pryor, C., Welch, D. L., & Fischer, P. 1993, *AJ*, 105, 510
- Mateo, M. L. 1998, *ARA&A*, 36, 435
- McConnachie, A. W., & Irwin, M. J. 2006, *MNRAS*, 365, 1263
- McGaugh, S. S. 2004, *ApJ*, 609, 652
- . 2005, *ApJ*, 632, 859
- McGaugh, S. S., de Blok, W. J. G., Schombert, J. M., Kuzio de Naray, R., & Kim, J. H. 2007, *ApJ*, 659, 149
- Milgrom, M. 1983, *ApJ*, 270, 365
- Minor, Q. E., Martinez, G., Bullock, J., Kaplinghat, M., & Trainor, R. 2010, *arXiv:1001.1160*
- Napolitano, N. R., Romanowsky, A. J., & Tortora, C. 2010, *arXiv:1003.1716*
- Olszewski, E. W., Pryor, C., & Armandroff, T. E. 1996, *AJ*, 111, 750
- Peñarrubia, J., Benson, A. J., Walker, M. G., Gilmore, G., McConnachie, A., & Mayer, L. 2010, *arXiv:1002.3376*
- Peñarrubia, J., Navarro, J. F., & McConnachie, A. W. 2008, *ApJ*, 673, 226
- Pryor, C., & Kormendy, J. 1990, *AJ*, 100, 127
- Sand, D. J., Olszewski, E. W., Willman, B., Zaritsky, D., Seth, A., Harris, J., Piatek, S., & Saha, A. 2009, *ApJ*, 704, 898
- Sanders, R. H., & McGaugh, S. S. 2002, *ARA&A*, 40, 263
- Simon, J. D., & Geha, M. 2007, *ApJ*, 670, 313
- Spano, M., Marcelin, M., Amram, P., Carignan, C., Epinat, B., & Hernandez, O. 2008, *MNRAS*, 383, 297
- Strigari, L. E., Bullock, J. S., Kaplinghat, M., Simon, J. D., Geha, M., Willman, B., & Walker, M. G. 2008, *Nature*, 454, 1096
- van der Hulst, J. M., Terlouw, J. P., Begeman, K. G., Zwitter, W., & Roelfsema, P. R. 1992, in *Astronomical Society of the Pacific Conference Series*, Vol. 25, *Astronomical Data Analysis Software and Systems I*, ed. D. M. Worrall, C. Biemesderfer, & J. Barnes, 131–+
- Walker, M. G., Mateo, M., Olszewski, E. W., Gnedin, O. Y., Wang, X., Sen, B., & Woodroffe, M. 2007, *ApJ*, 667, L53
- Walker, M. G., Mateo, M., Olszewski, E. W., Peñarrubia, J., Wyn Evans, N., & Gilmore, G. 2009, *ApJ*, 704, 1274
- Wilkinson, M. I., Kleyna, J., Evans, N. W., & Gilmore, G. 2002, *MNRAS*, 330, 778
- Wolf, J., Martinez, G. D., Bullock, J. S., Kaplinghat, M., Geha, M., Munoz, R. R., Simon, J. D., & Avedo, F. F. 2009, *arXiv:0908.2995*